

Hands-on Session on Data Analysis SIR code

Salvo Guglielmino - INAF Catania

"The different spatio-temporal scales of the solar magnetism"
12 April 2022

I'm deeply indebted to ...

- Luis Bellot Rubio
 - THE ART OF STOKES INVERSIONS
 - Lesson at the "Solar Magnetic Fields: Modeling and Measuring Techniques School", 25-28 May 2015, Granada
 - Lesson at the "Spectropolarimetry and Diagnostic Techniques School", 24-Sep / 05-Oct 2018, Estes Park
- Basilio Ruiz Cobo / José Carlos del Toro Iniesta
 - Inversion of the radiative transfer equation for polarized light (2016), Living Rev. Sol. Phys. 13, 4

Inversion codes: M-E vs full radiative transfer

Milne-Eddington

- Results
 - robust
 - easy to interpret
- Analytical Stokes profiles
- Fast inversion
- Asymmetric profiles
- Poor thermal information
- No multi-component atmospheres in the resolution element

Full radiative transfer

- Numerical solution of the RTE
- © Reliable thermal information
- \odot Stratification of the physical parameter with depth (τ)
- Deal with asymmetries
- Allow multi-component atmospheres
- © Computationally demanding: slow inversion
- Difficult to interpret

Asymmetric Stokes profiles

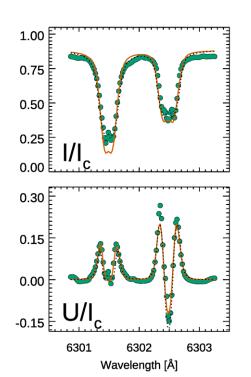
Amplitude asymmetry/
Multi-lobed Stokes profiles

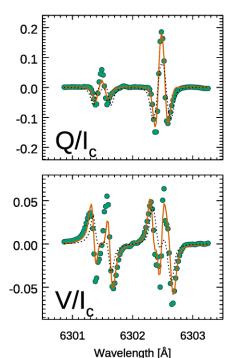
Area asymmetry

Different magnetic atmospheres coexisting in resolution element

Gradients/discontinuities of B and v_{LOS} along LOS

Auer & Heasley (1978)





The area asymmetry gives information on the **height variation** of atmospheric parameters

SIR

STOKES INVERSION BASED ON RESPONSE FUNCTIONS
Ruiz Cobo & del Toro Iniesta (1992), ApJ 398, 375

The SIR code

 Provides a numerical solution of the full radiative transfer equation in presence of magnetic field

$$\frac{d}{d\tau} \begin{pmatrix} I \\ Q \\ U \\ V \end{pmatrix} = \begin{pmatrix} \eta_{\mathrm{l}} & \eta_{\mathrm{Q}} & \eta_{\mathrm{U}} & \eta_{\mathrm{V}} \\ \eta_{\mathrm{Q}} & \eta_{\mathrm{l}} & \rho_{\mathrm{V}} & -\rho_{\mathrm{U}} \\ \eta_{\mathrm{U}} & -\rho_{\mathrm{V}} & \eta_{\mathrm{l}} & \rho_{\mathrm{Q}} \\ \eta_{\mathrm{V}} & \rho_{\mathrm{U}} & -\rho_{\mathrm{Q}} & \eta_{\mathrm{l}} \end{pmatrix} \begin{pmatrix} I - S \\ Q \\ U \\ V \end{pmatrix}$$
 (Unno 1956; Rachkovsky 1962)

- ► Local Thermodynamic Equilibrium (LTE) assumption
- Handles Stokes parameters and works in two modes:
 - synthesis mode / inversion mode
- Allows considering either one or two magnetic atmospheres
- Retrieves the thermal, dynamical and magnetic structure of the atmosphere as a function of τ_{500}

The SIR code: a few caveats

- Here, we will discuss the "classical" approach to SIR, devoted to a single-pixel inversion
- Other ad hoc versions of the SIR code exist
- Note that, very recently, major updates have been made, including a Python wrapper for the SIR
 - this allows the user to upload a .fits file containing the observed data and to split the field-of-view among the different CPUs, working in parallel
- ► Furthermore, SIR is now combined with the RH code to treat non-LTE problems: DeSIRe code
 - see Ruiz Cobo et al. (2022), A&A in press
- You can benefit from the free online courses by Carlos Quintero Noda (IAC)
 - youtube.com/channel/UCTR42H92ZJMjv_snYK9Xfmw

Response functions

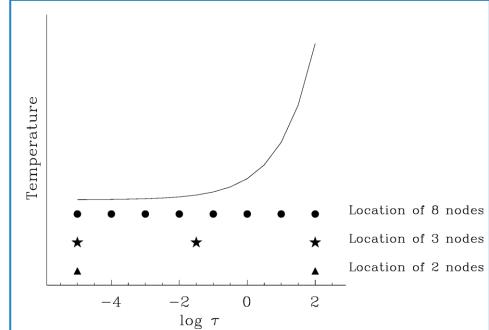
- Can be considered the response of the observed spectrum to the modifications of an atmosphere model
- When changing physical conditions along the optical depth τ , this perturbation results in a modification of the emergent Stokes spectrum

$$\delta I(\lambda) = \int_0^\infty R(\lambda, \, \tau) \delta x(\tau) d\tau$$

- At first-order perturbative analysis of the RTE:
 - $\delta x(\tau)$ is a perturbation of a single physical parameter $x(\tau)$
 - it propagates resulting in a modification $\delta I(\lambda)$ of the emergent spectrum $I(\lambda)$
 - the quantity $R(\lambda, \tau)$ is a vector containing the response function for each Stokes parameter

The nodes

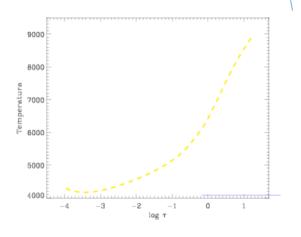
- ► From an operational point of view, equivalent perturbations (RFs) of a single physical quantity are calculated at certain locations (grid points) the nodes with a cubic-splines interpolation between them
- The optical depth τ is equi-spaced on a logarithmic scale
- Nodes are also equi-spaced
- The number of nodes can be set by the user or automatically optimized by the code
 - no. = 0 -> not inverted
 - no. = 1 -> constant quantity
 - no. = 2 -> linear interpolation



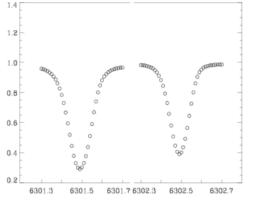
Procedure

- The code iteratively changes each physical quantity, assuming initial constant values for $B(\tau)$, $\gamma(\tau)$ and $\varphi(\tau)$ set by the user or deduced from the weak field approximation
- Iteration is repeated to reach convergence, by minimizing the χ^2 merit function containing the differences between the observed and synthetic spectral data

$$\chi^2 \equiv \frac{1}{v} \sum_{k=1}^{4} \sum_{i=1}^{M} \left[I_k^{\text{obs}}(\lambda_i) - I_k^{\text{syn}}(\lambda_i) \right]^2$$

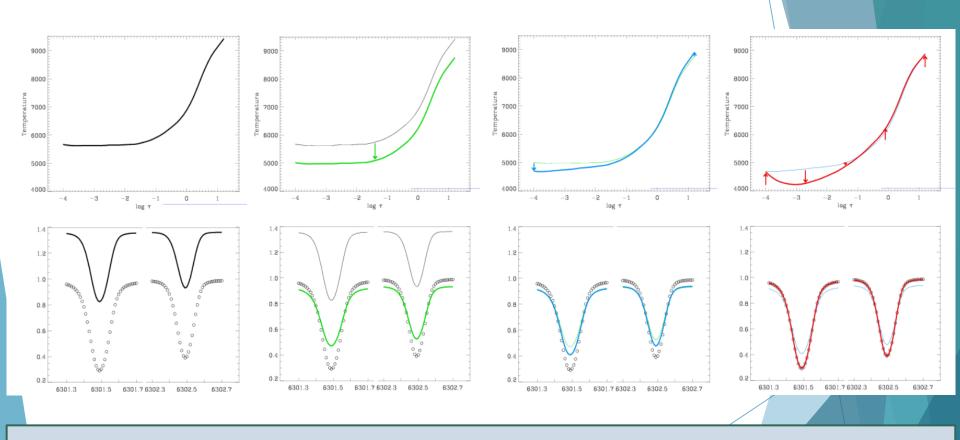


The "real" SUN



Synthetic "observations"

Iterations



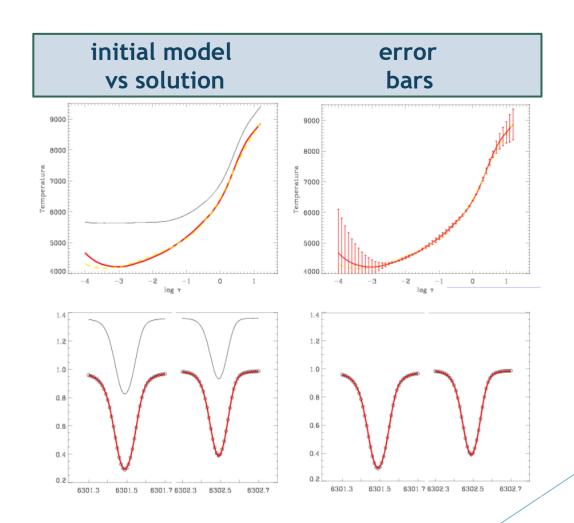
1 node

initial "guess" model

2 nodes

5 nodes

Output



https://github.com/BasilioRuiz/SIR-code

REPOSITORY FOR THE SIR CODE originally developed in FORTRAN

Running SIR

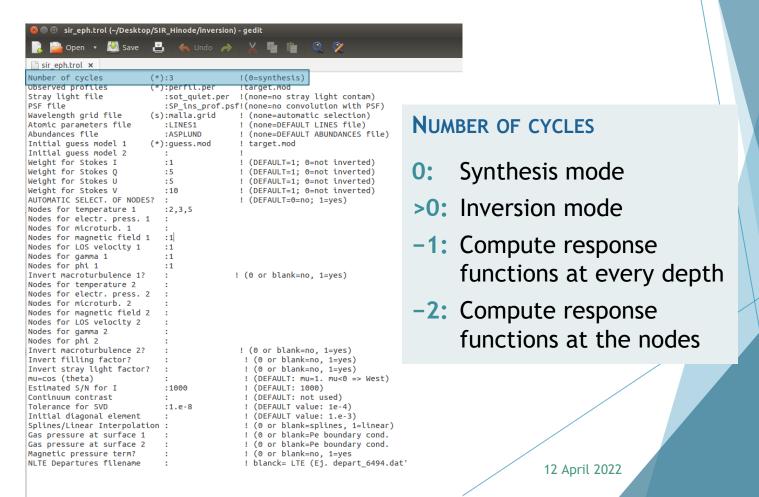
A file [].trol controls the code

```
sir_eph.trol (~/Desktop/SIR_Hinode/inversion) - gedit
     📔 Open 🔻 💹 Save 📲 👆 Undo 🧀
sir_eph.trol ×
Number of cycles
                          (*):3
                                             !(0=synthesis)
Observed profiles
                          (*):perfil.per
                                             !target.mod
Stray light file
                             :sot quiet.per !(none=no stray light contam)
PSF file
                             :SP_ins_prof.psf!(none=no convolution with PSF)
Wavelength grid file
                          (s):malla.grid
                                             ! (none=automatic selection)
Atomic parameters file
                             :LINES1
                                             ! (none=DEFAULT LINES file)
Abundances file
                             :ASPLUND
                                             ! (none=DEFAULT ABUNDANCES file)
Initial guess model 1
                          (*):guess.mod
                                             ! target.mod
Initial guess model 2
Weight for Stokes I
                                             ! (DEFAULT=1; 0=not inverted)
                             : 1
                                             ! (DEFAULT=1; 0=not inverted)
Weight for Stokes 0
                             : 5
Weight for Stokes U
                             : 5
                                             ! (DEFAULT=1; 0=not inverted)
Weight for Stokes V
                             :10
                                             ! (DEFAULT=1; 0=not inverted)
AUTOMATIC SELECT. OF NODES?
                                             ! (DEFAULT=0=no; 1=yes)
Nodes for temperature 1
                             :2,3,5
Nodes for electr. press. 1
Nodes for microturb. 1
Nodes for magnetic field 1
Nodes for LOS velocity 1
Nodes for gamma 1
Nodes for phi 1
Invert macroturbulence 1?
                                           ! (0 or blank=no, 1=yes)
Nodes for temperature 2
Nodes for electr. press. 2
Nodes for microturb. 2
Nodes for magnetic field 2
Nodes for LOS velocity 2
Nodes for gamma 2
Nodes for phi 2
                                             ! (0 or blank=no, 1=yes)
Invert macroturbulence 2?
Invert filling factor?
                                             ! (0 or blank=no, 1=yes)
Invert stray light factor?
                                             ! (0 or blank=no, 1=yes)
                                             ! (DEFAULT: mu=1. mu<0 => West)
mu=cos (theta)
Estimated S/N for I
                             :1000
                                             ! (DEFAULT: 1000)
Continuum contrast
                                              ! (DEFAULT: not used)
Tolerance for SVD
                                             ! (DEFAULT value: 1e-4)
                             :1.e-8
Initial diagonal element
                                              ! (DEFAULT value: 1.e-3)
Splines/Linear Interpolation :
                                              ! (0 or blank=splines, 1=linear)
Gas pressure at surface 1 :
                                              ! (0 or blank=Pe boundary cond.
Gas pressure at surface 2
                                              ! (0 or blank=Pe boundary cond.
Magnetic pressure term?
                                              ! (0 or blank=no, 1=yes
NLTE Departures filename
                                              ! blanck= LTE (Ej. depart_6494.dat'
```

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Running SIR: cycles

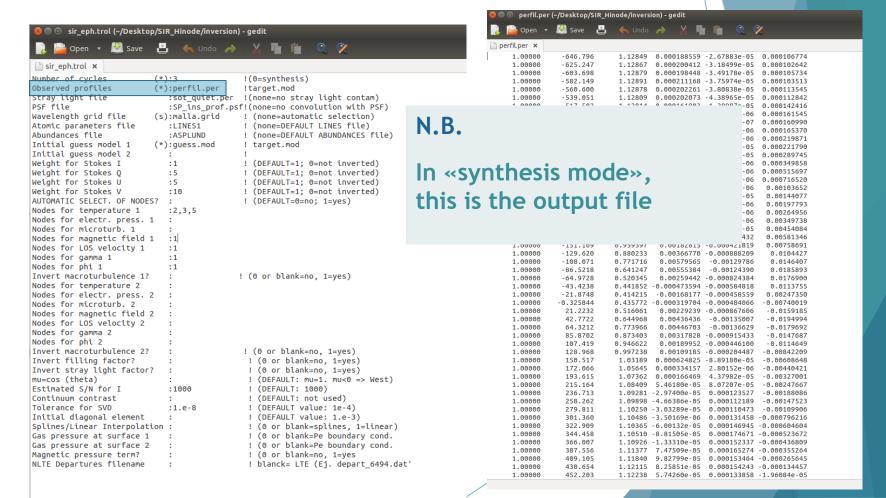
A file [].trol controls the code



Running SIR: observed data

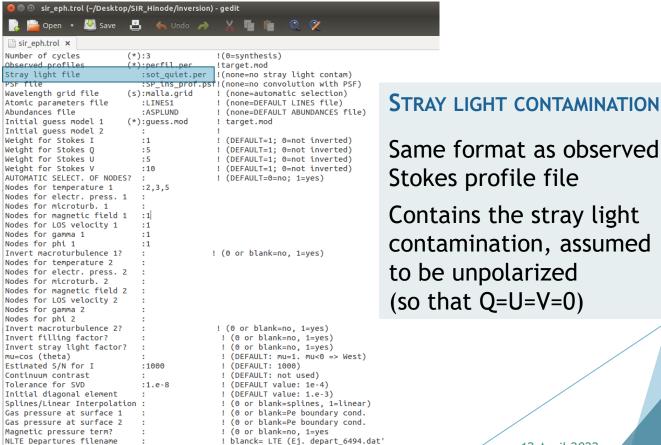
► A file [].trol controls the code

Line $\Delta\lambda$ index [mA] I/I_{qs} Q/I_{qs} U/I_{qs} V/I_{qs}



Running SIR: stray light

► A file [].trol controls the code



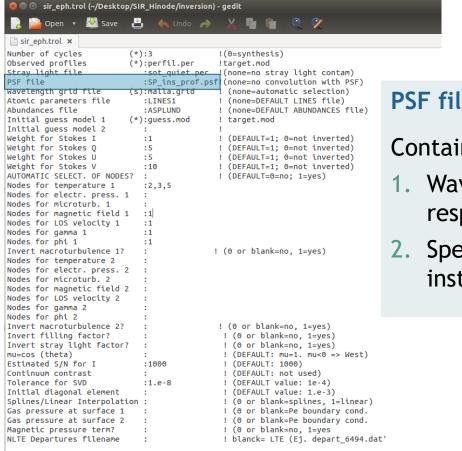
Stray-light treatment

- Stray-light in 1C inversions:

 - Accounts for both stray light and/or magnetic filling factor
- Stray-light in 2C inversions:
 - It is NOT equivalent to a magnetic filling factor
 - SIR has two free parameters: α and f
 - $I_{obs} = (1-\alpha) [f I_1 + (1-f) I_2] + \alpha I_{stray}$
- Global vs local stray-light profile
 - Classical treatment: global stray-light profile (average over FOV)
 - local stray-light profile to account for telescope diffraction (e.g. Orozco Suárez et al., 2007)

Running SIR: PSF deconvolution

A file [].trol controls the code



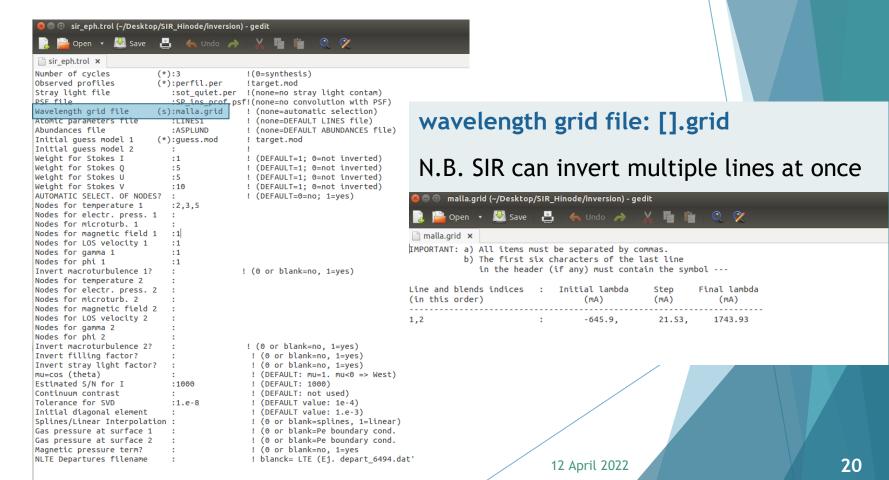
PSF file: [].psf

Contains two columns:

- 1. Wavelength in mA with respect to center of the line
- Spectral PSF of the instrument

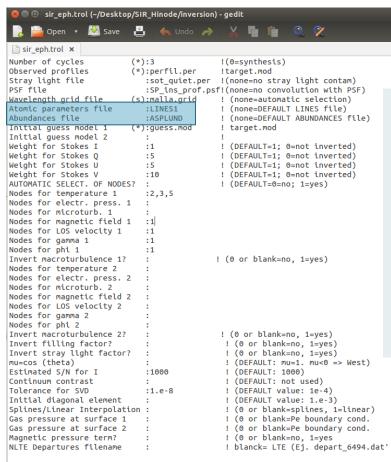
Running SIR: wavelength grid

A file [].trol controls the code



Running SIR: wavelength grid

► A file [].trol controls the code



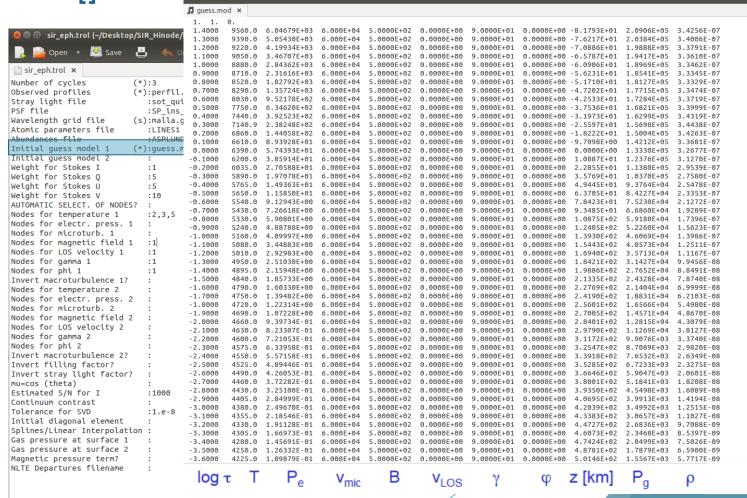
These files contain:

- For a given line, identified with the «line index» ion λ E χ log gf transition collisional parameters
- Abundances of the different atomic species

Running SIR: model (I)

atmospheric model file: [].mod

► A file [].trol cor per value Save 4 on the cord process.mod (-/Desktop/SIR_Hinode/Inversion) - gedit

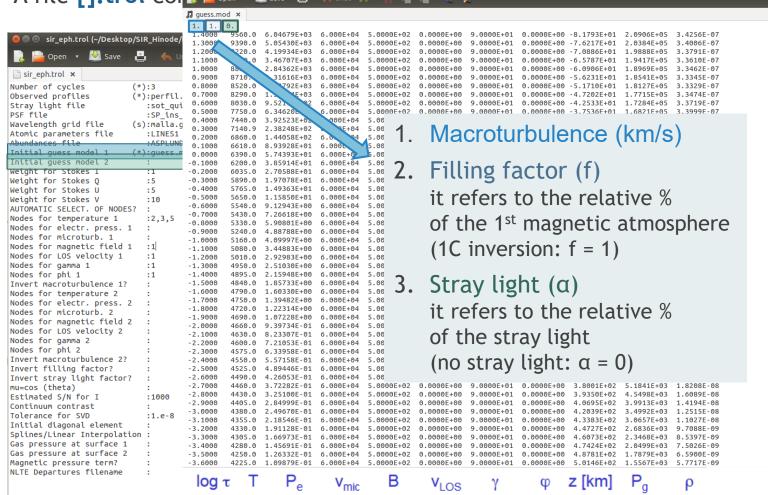


Running SIR: model (II)

atmospheric model file: [].mod

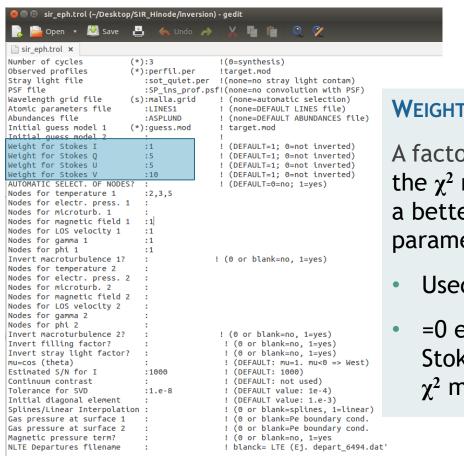
A file [].trol cor ☐ guess.mod (-/Desktop/SIR_Hinode/Inversion) - gedit

| A file [].trol cor ☐ ☐ Open • ☐ Save ☐ ← Undo → ※ ☐



Running SIR: weights for Stokes

A file [].trol controls the code



WEIGHTS

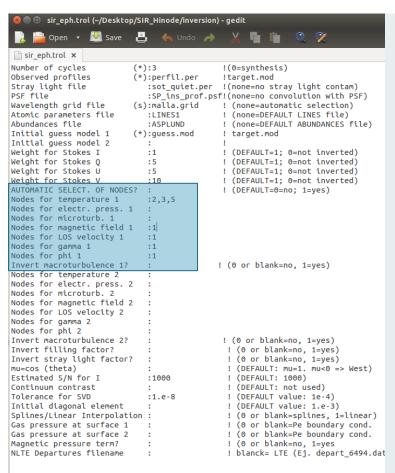
A factor that can be added into the χ^2 merit function to force a better fit for a given Stokes parameter (Q, U, V)

- Used for peculiar profiles
- =0 excludes the fit to the Stokes parameter from the χ^2 merit computation

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Running SIR: choice of nodes

► A file [].trol controls the code



NUMBER OF NODES

Each column corresponds to a cycle

E.g. assume we have 3 cycles with

- nodes for T: 2,3,5
- nodes for B/v_{LOS}: 1,2
- In the first cycle, 2 nodes for T and 1 for B/v_{LOS} will be used
- In the second cycle, 3 for T and 2 for B/v_{LOS} will be used
- In the third cycle, 5 for T and again 2 for B/v_{LOS} will be used

0 nodes in electron pressure -> HE

no. of nodes * means any value
(only for automatic selection)

Running SIR: other settings

A file [].trol controls the code



OTHER INVERSION SETTINGS

- S/N ratio are related to the acceptable quality of the fit, compared to the instrumental noise
- SVD tolerance indicates the convergence threshold in the χ^2 merit function for the iterations

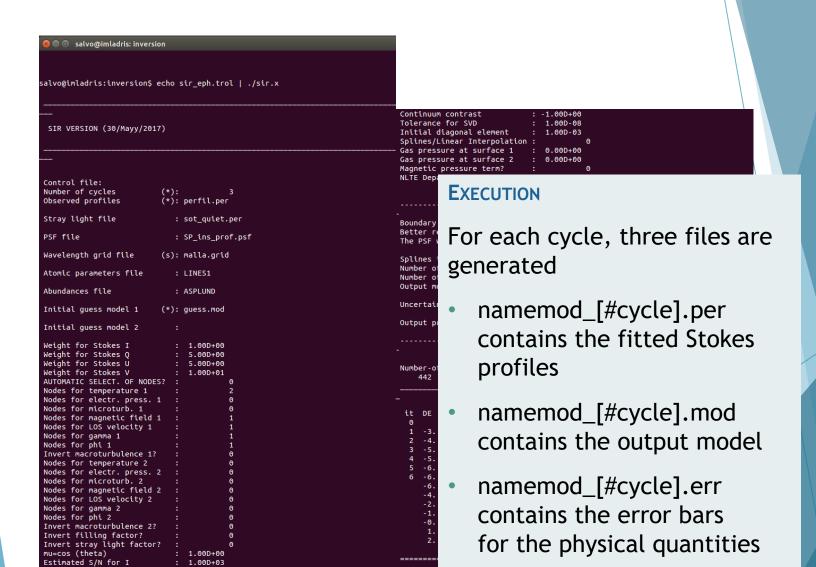
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Executing SIR

```
❷ □ salvo@imladris:inversion
salvo@imladris:inversion$ echo sir_eph.trol | ./sir.x□
```

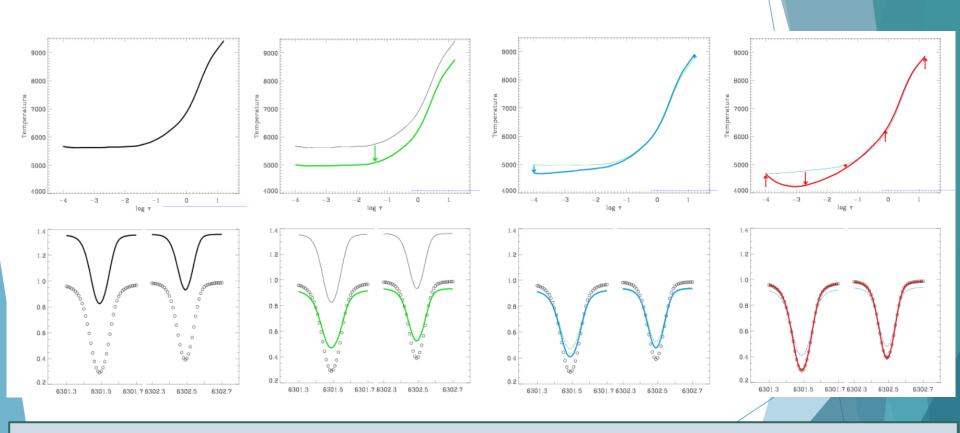
echo sir.trol | ./sir.x

Executing SIR



Executing SIR

3 cycles



guess.mod profile.per

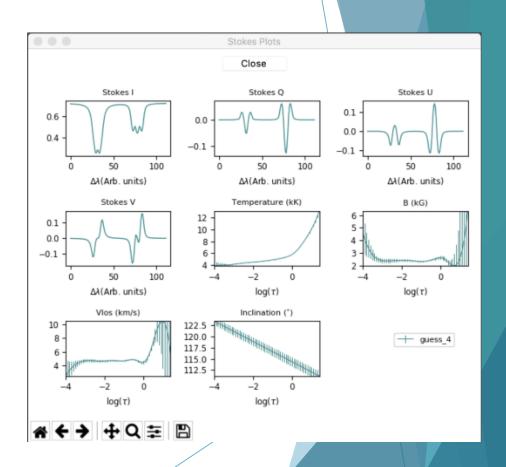
guess_1.mod
guess_1.per

guess_2.mod
guess_2.per

guess_3.mod
guess_3.per

Visualizing SIR results

- ► IDL GRAPHICS.PRO
 - included in the SIR standard distribution
 - [path]/idl/graphics.pro
- Python SIRGUI.PY
 - https://github.com/rcenteno/ SIR_GUI



Milne-Eddington-like inversion

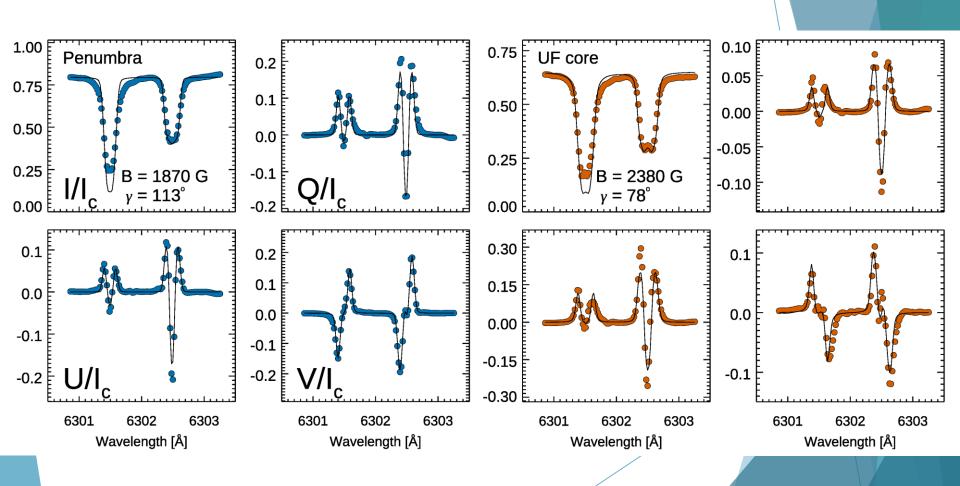
- ► An inversion scheme, i.e., a set of cycles and nodes that has only 1 node for all the physical parameters except for T, which has 2+ nodes for all the cycles, can be treated as a Milne-Eddington-like inversion
- Main advantages:
 - "fast" inversion (with respect to other full RTE schemes)
 - realistic thermal stratification
 - comparable with ME results
 - easy to interpret
- Contra:
 - asymmetries cannot be treated
 - much slower than ME approach

Milne-Eddington-like inversion

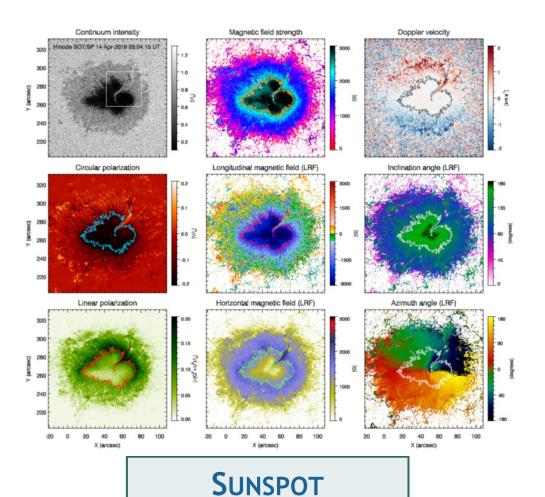
```
sir_eph.trol (~/Desktop/SIR_Hinode/inversion) - gedit
     ๊ Open 🔻 💹 Save 🖺
 sir_eph.trol x
Number of cycles
                                              !(0=svnthesis)
Observed profiles
                          (*):perfil.per
                                              !target.mod
Stray light file
                              :sot quiet.per !(none=no stray light contam)
PSF file
                              :SP ins prof.psf!(none=no convolution with PSF)
Wavelength grid file
                          (s):malla.grid
                                              ! (none=automatic selection)
Atomic parameters file
                              :LINES1
                                              ! (none=DEFAULT LINES file)
Abundances file
                              :ASPLUND
                                              ! (none=DEFAULT ABUNDANCES file)
Initial guess model 1
                                              ! target.mod
                          (*):quess.mod
Initial guess model 2
Weight for Stokes I
                                              ! (DEFAULT=1; 0=not inverted)
                              : 1
Weight for Stokes Q
                              :5
                                              ! (DEFAULT=1; 0=not inverted)
Weight for Stokes U
                              :5
                                              ! (DEFAULT=1; 0=not inverted)
Weight for Stokes V
                              :10
                                              ! (DEFAULT=1; 0=not inverted)
AUTOMATIC SELECT. OF NODES?
                                              ! (DEFAULT=0=no; 1=yes)
Nodes for temperature 1
Nodes for electr. press. 1
Nodes for microturb, 1
Nodes for magnetic field 1
Nodes for LOS velocity 1
                             :1
Nodes for gamma 1
                             :1
Nodes for phi 1
Invert macroturbulence 1?
                                             ! (0 or blank=no, 1=yes)
Nodes for temperature 2
Nodes for electr. press. 2
Nodes for microturb. 2
Nodes for magnetic field 2
Nodes for LOS velocity 2
Nodes for gamma 2
Nodes for phi 2
                                              ! (0 or blank=no, 1=ves)
Invert macroturbulence 2?
Invert filling factor?
                                               ! (0 or blank=no, 1=ves)
                                               ! (0 or blank=no, 1=yes)
Invert stray light factor?
mu=cos (theta)
                                               ! (DEFAULT: mu=1. mu<0 => West)
Estimated S/N for I
                             :1000
                                               ! (DEFAULT: 1000)
Continuum contrast
                                               ! (DEFAULT: not used)
Tolerance for SVD
                             :1.e-8
                                               ! (DEFAULT value: 1e-4)
                                               ! (DEFAULT value: 1.e-3)
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Splines/Linear Interpolation :
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Gas pressure at surface 1
                                               ! (0 or blank=Pe boundary cond.
Gas pressure at surface 2
                                               ! (0 or blank=Pe boundary cond.
                                               ! (0 or blank=no, 1=yes
Magnetic pressure term?
                                               ! blanck= LTE (Ej. depart_6494.dat'
NLTE Departures filename
```

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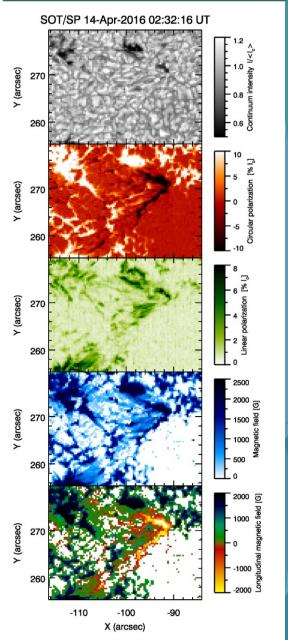
Milne-Eddington-like inversion: examples



Milne-Eddington-like inversion: contexts

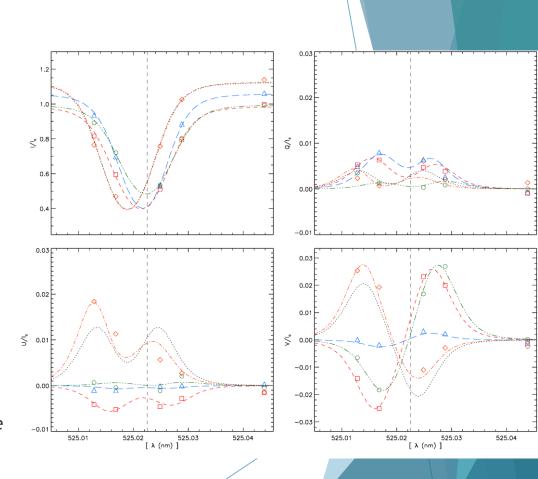


EMERGING FLUX REGION

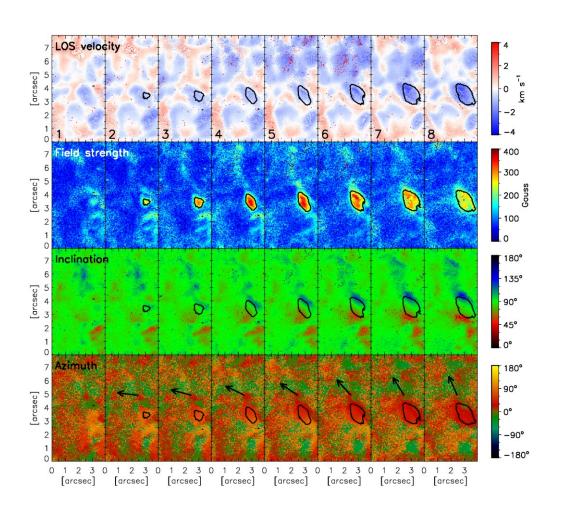


Using linear gradients

- Observed (symbols) I, Q, U, V profiles, for a sample of IMaX spectra, with the corresponding fits (lines) obtained with the SIR code
 - blue (triangles/long dashed): horizontal fields, no gradients
 - red (squares/dashed): vertical fields, no gradients
 - green (circle/dot-dot-dot dashed): asymmetry in Stokes V with weak L_s signal, with gradients
 - orange (diamonds/dot dashed): asymmetry in the blue and red lobes of Stokes U and V, with gradients
- The dotted line represents the fit without gradients to the plotted orange spectrum (for Stokes I the dotted line falls on top of the orange one)
- The dashed vertical line indicates the nominal line center λ

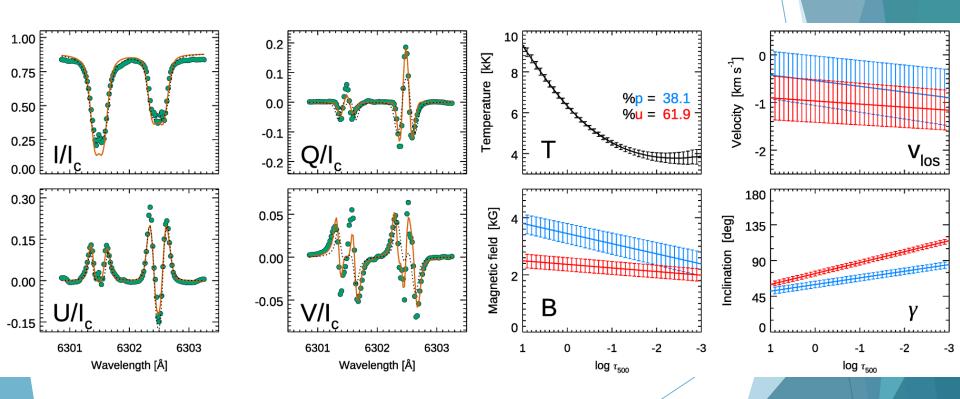


Using linear gradients

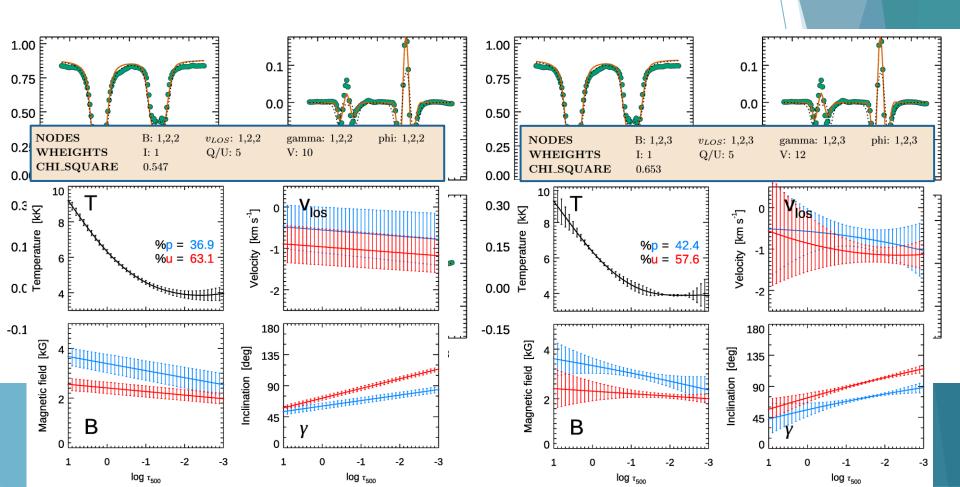


EVALUATION AT LOG $\tau = -2$

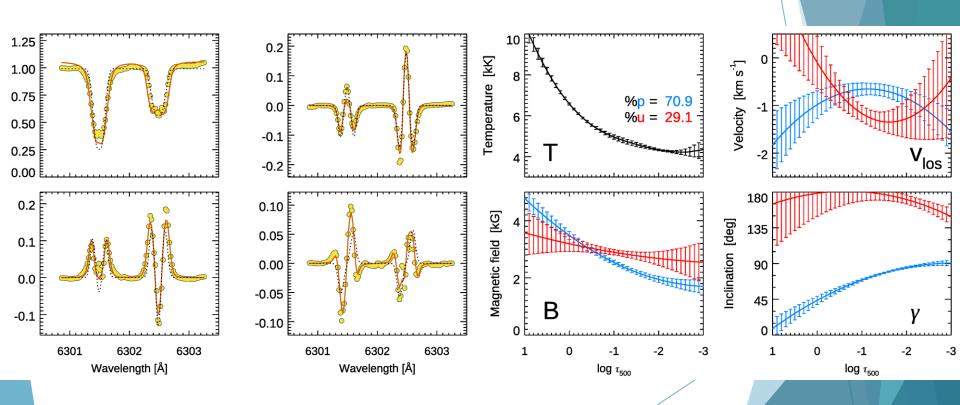
Two-component inversions: gradients are enough



Two-component inversions: comparing results



Two-component inversions: a more complex situation



Conclusions

- ► Full RTE inversions are able to give significant physical information, although they are computationally costly
- Examine the fits: are they reasonably good with M-E?
- SIR is a flexible and powerful inversion code in LTE
- Be careful with the choice of atmospheric model!
 The interpretation is not straightforward...
 - always start with a M-E-like inversion
 - keep the number of free parameters (nodes) small
 - ask yourself if the retrieved model atmosphere makes sense
- Uniqueness problems
 - results may change if the physical model is changed